

# Probing galaxy evolution in groups with Chandra and XMM



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# Why study (galaxy evolution in) groups?

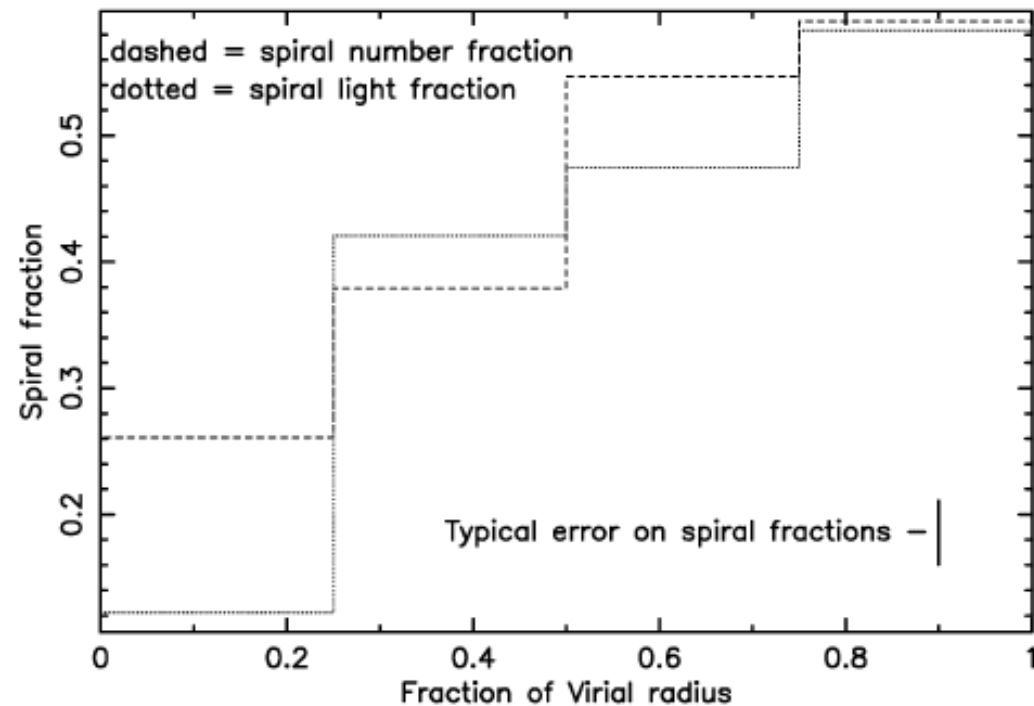
Most galaxies

- not isolated
- affected by environment.

Clusters:  
rare + “evolved”

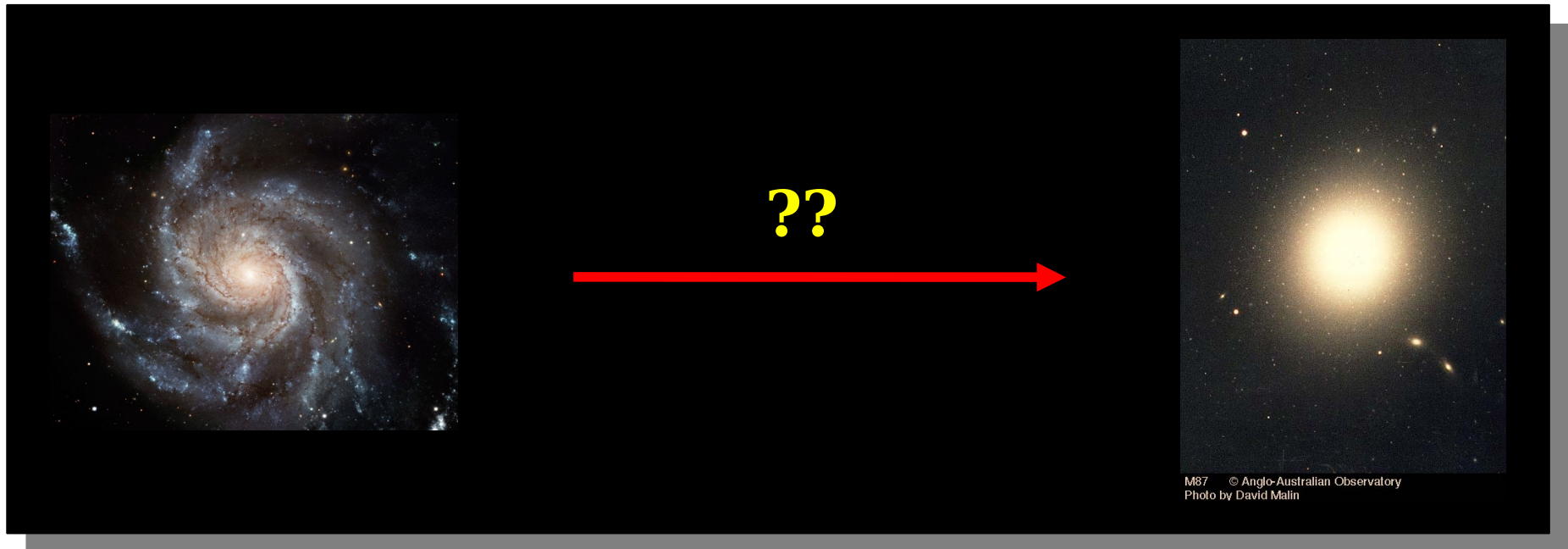
Groups:  
Common + “evolving”

24 X-ray bright groups:



Helsdon & Ponman 2003

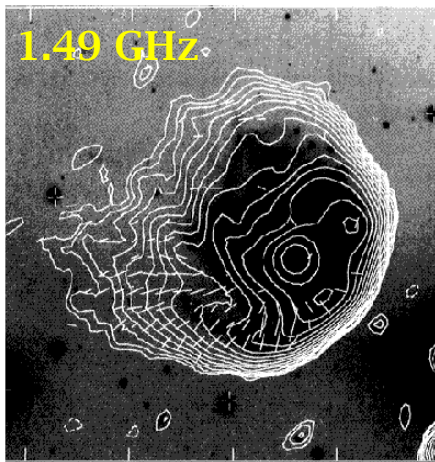
# Galaxy transformations



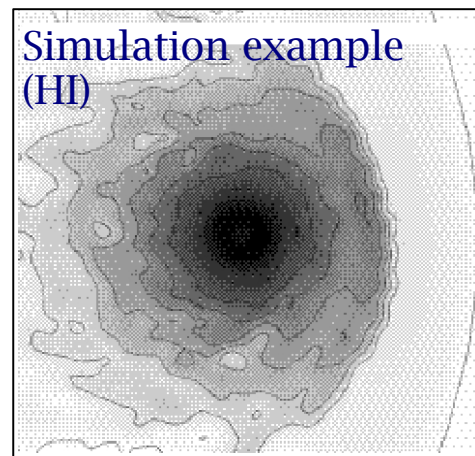
- Galaxy-galaxy interactions
- Galaxy-gas interactions: Ram pressure  $\mathbf{P} = \rho_{\text{gas}} \mathbf{v}_{\text{gal}}^2$

# NGC 2276 – a transforming spiral

Starbursting ( $\sim 5 M_{\odot}/\text{yr}$ ) spiral embedded in  $T \approx 10^7 \text{ K}$  group gas.



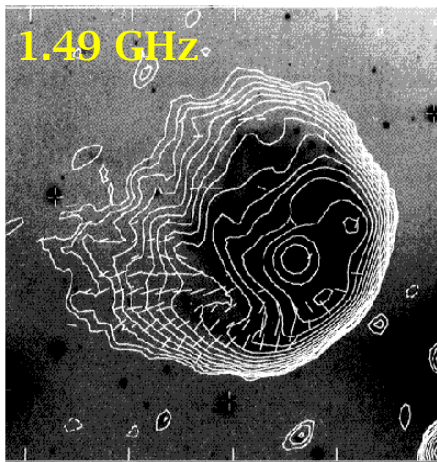
Hummel & Beck 1995



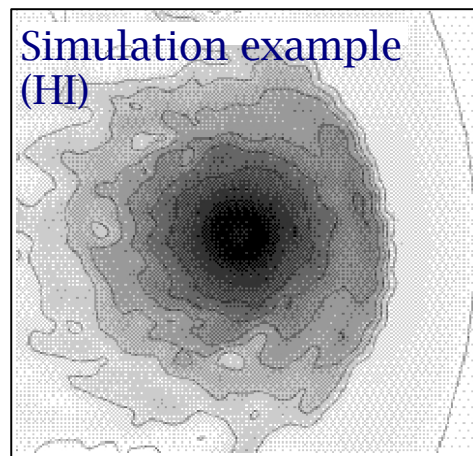
Quilis et al. 2000

# NGC 2276 – a transforming spiral

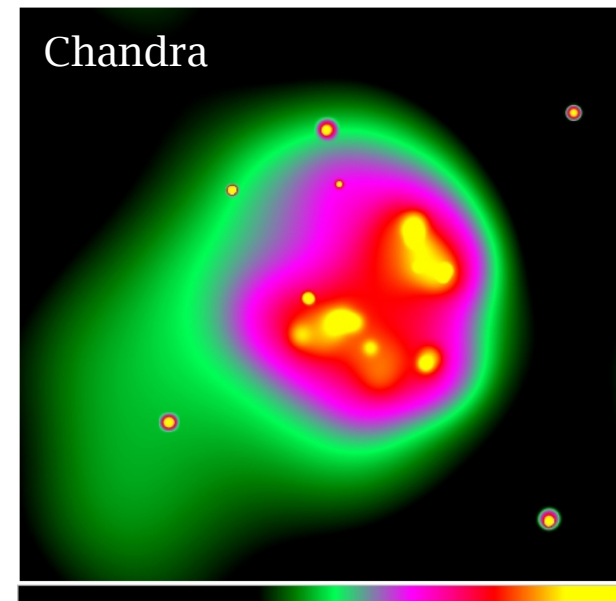
Starbursting ( $\sim 5 M_{\odot}/\text{yr}$ ) spiral embedded in  $T \approx 10^7 \text{ K}$  group gas.



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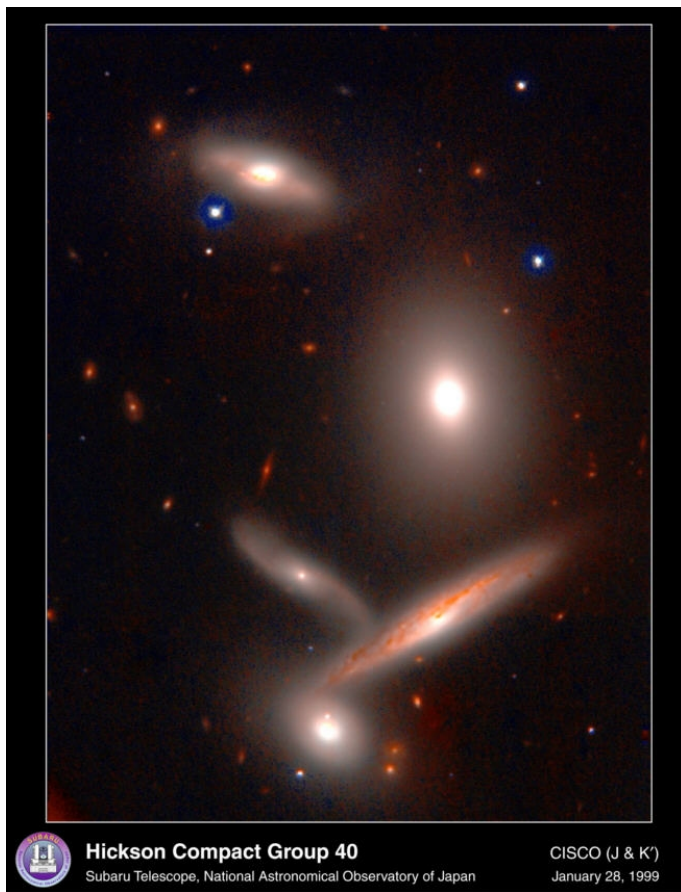
- $v_{\text{gal}} \approx 850 \text{ km/s}$
- Current mass-loss rate of gas:  $\sim 5 M_{\odot}/\text{yr}$ .
- Gas supply lost in 1-2 Gyr

→ S0 ?



# HI-deficient compact groups

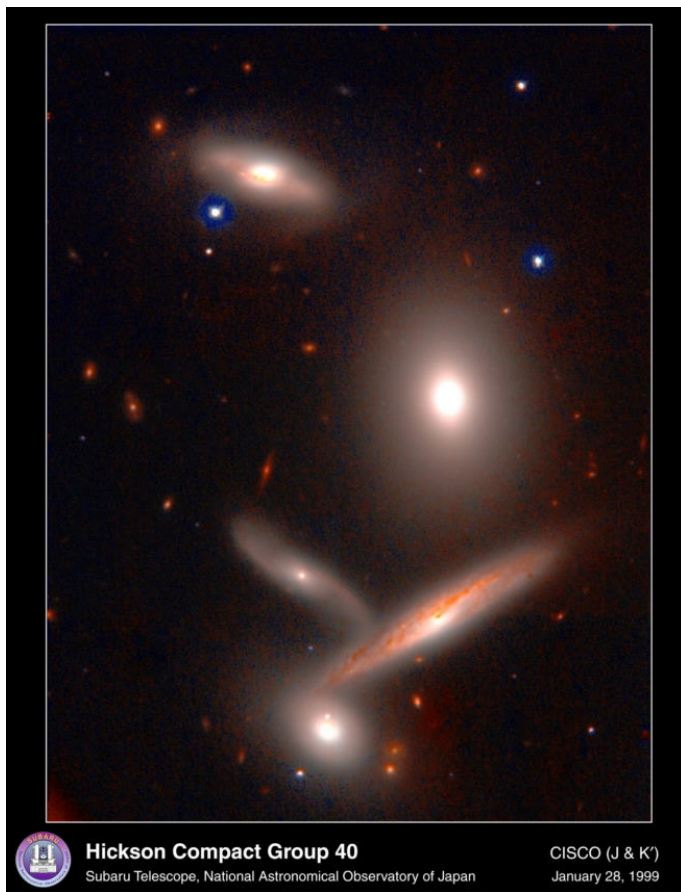
Compact groups (Hickson 1982): Deficient in HI.



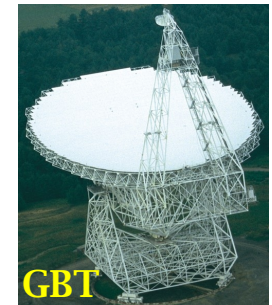
HI-deficiency related to X-ray  
emission from group gas?  
(Verdes-Montenegro+ 01)

# HI-deficient compact groups

Compact groups (Hickson 1982): Deficient in HI.

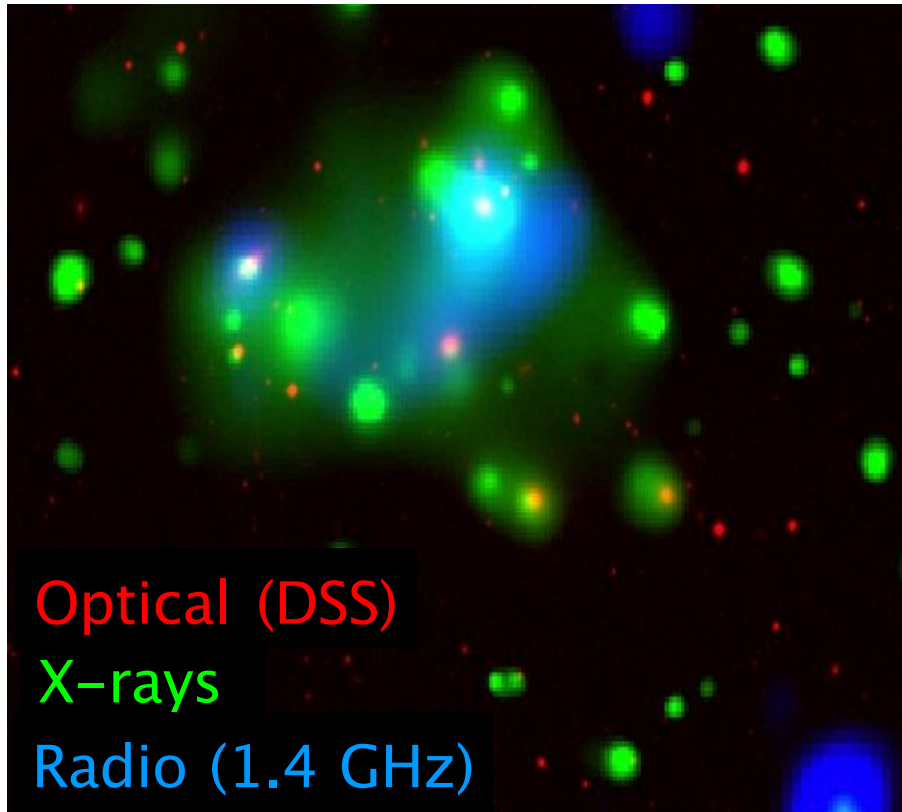


HI-deficiency related to X-ray emission from group gas?  
(Verdes-Montenegro+ 01)



# Two examples: HI vs. hot gas

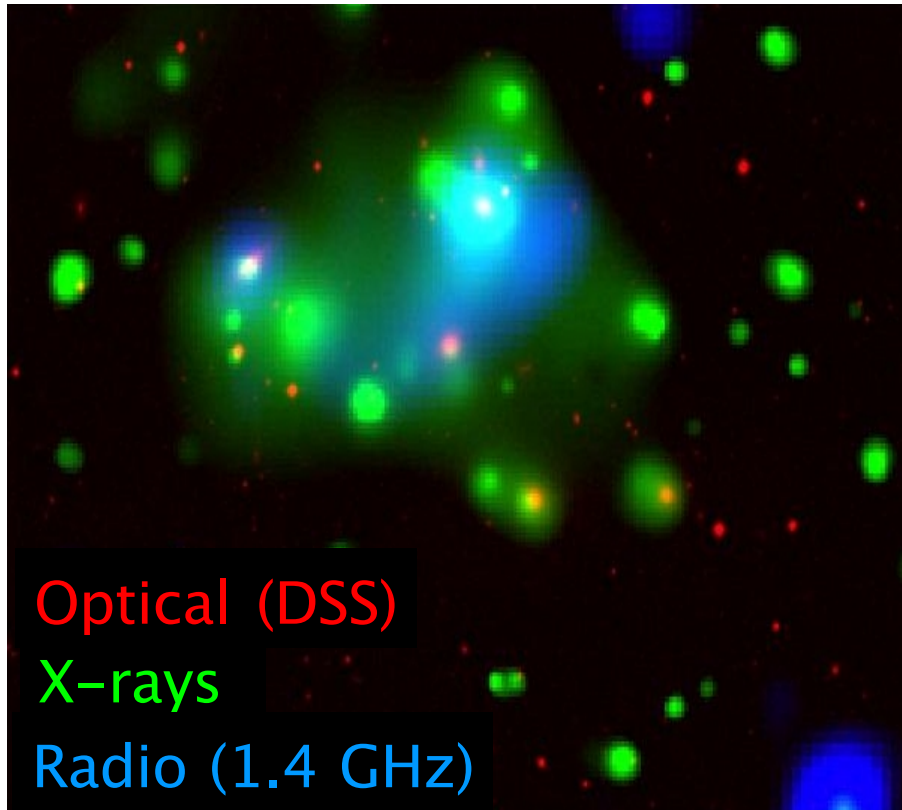
HCG 15



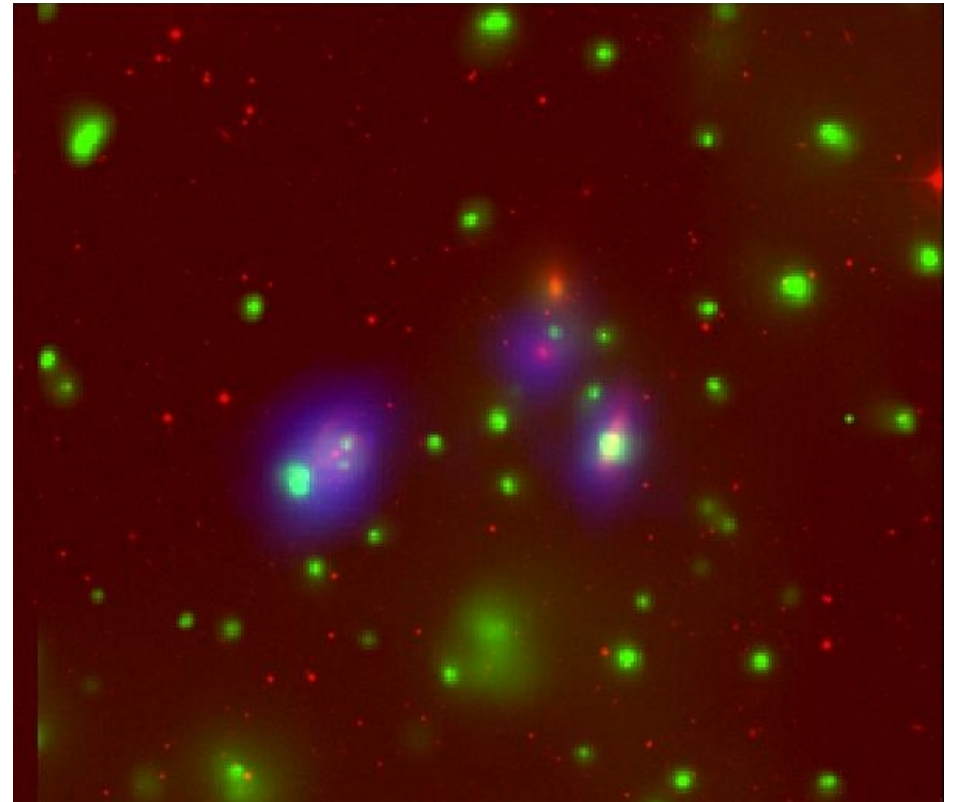


# Two examples: HI vs. hot gas

HCG 15

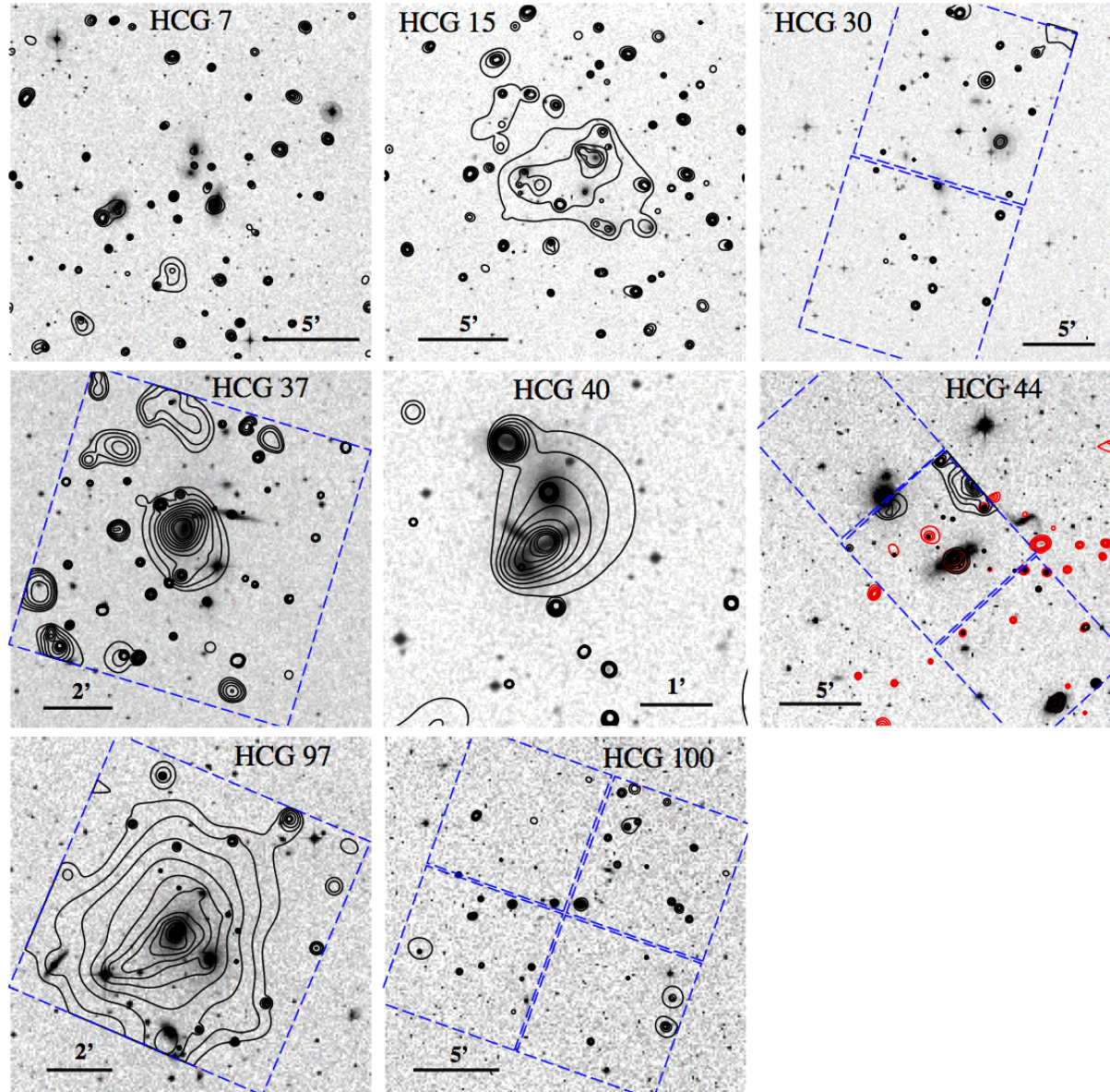


HCG 7

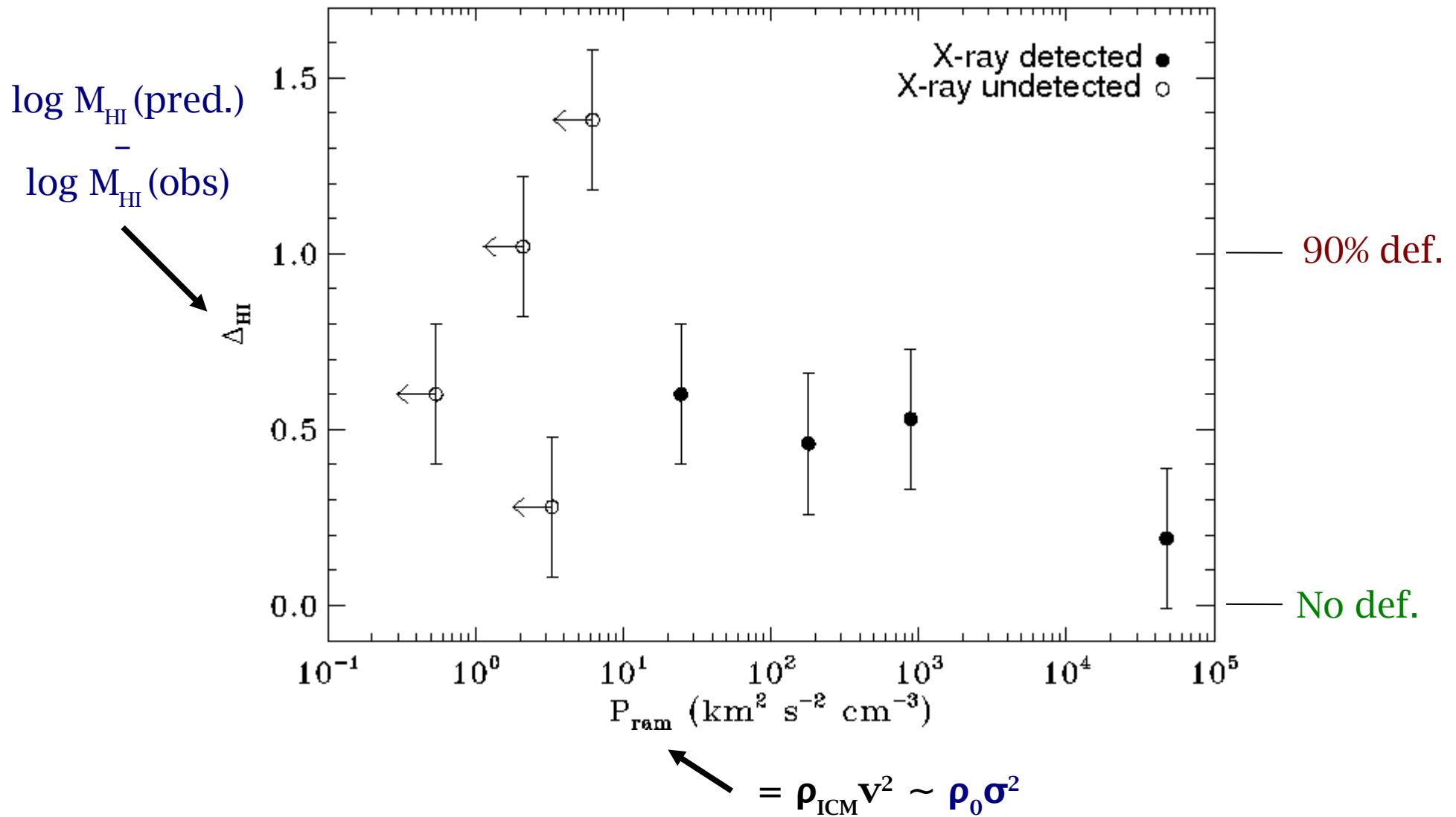


# X-ray gallery of Hickson groups

Chandra/XMM  
0.3-2 keV  
contours over  
DSS images



# HI deficiency: Caused by ram pressure?





# Simple modeling of ICM interactions

Ram pressure stripping:  $F_{\text{grav}}/D^2 < \rho_{\text{ICM}} v^2$



- DM halo
  - bulge
- gas+stellar disks



# Simple modeling of ICM interactions

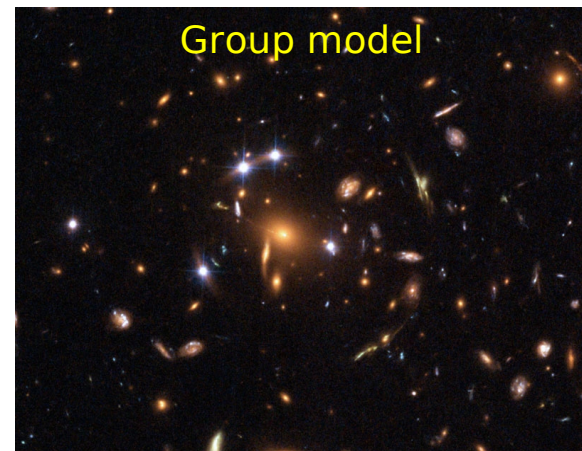
Ram pressure stripping:  $F_{\text{grav}}/D^2 < \rho_{\text{ICM}} v^2$



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Radial  
free-fall

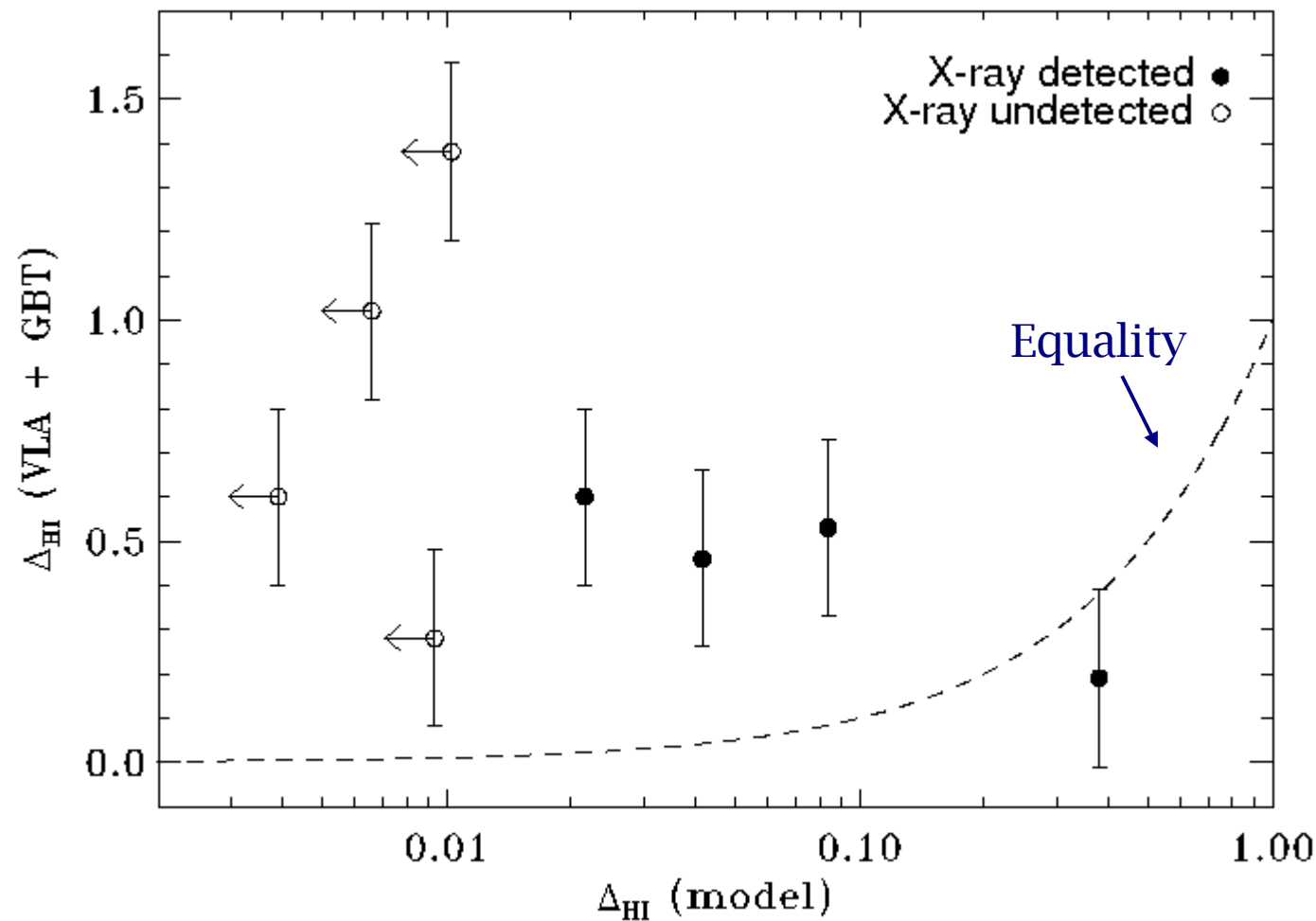


- NFW potential
- Measured gas profile (eventually!...)

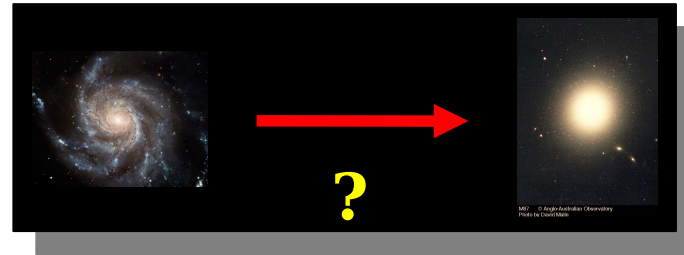
$$\Sigma_g \left( \frac{\partial \Phi_b}{\partial z} + \frac{\partial \Phi_h}{\partial z} + \frac{\partial \Phi_g}{\partial z} + \frac{\partial \Phi_*}{\partial z} \right) < \rho_1 (v_1 \sin \xi)^2$$

# Mass loss from ram pressure

From predicted mass loss:



# Conclusions & future work...



- HI removal in groups:  
Ram pressure stripping not the whole story.

- Destruction of HI:  
Not due to ICM heating in all cases.

- *Need to consider:*  
Realistic orbits, DM halo (tidal) stripping,  
viscous stripping (cf. N2276)...